


Experimental constraint on stellar electron-capture rates from the $^{88}\text{Sr}(t, ^3\text{He} + \gamma) ^{88}\text{Rb}$ reaction at 115 MeV/u

J. C. Zamora ^{1,2}, R. G. T. Zegers,^{1,2,3} Sam M. Austin,^{1,2} D. Bazin,¹ B. A. Brown,^{1,2,3} P. C. Bender,¹ H. L. Crawford,⁴ J. Engel,⁵ A. Falduto,^{6,2} A. Gade,^{1,2,3} P. Gastis,^{6,2} B. Gao,^{1,2} T. Ginter,¹ C. J. Guess,⁷ S. Lipschutz,^{1,2,3} B. Longfellow,^{1,3} A. O. Macchiavelli,⁴ K. Miki,⁸ E. Ney,⁵ S. Noji,^{1,2} J. Pereira,^{1,2} J. Schmitt,^{1,2,3} C. Sullivan,^{1,2,3} R. Titus,^{1,2,3} and D. Weisshaar¹

¹National Superconducting Cyclotron Laboratory, Michigan State University, East Lansing, Michigan 48824, USA

²Joint Institute for Nuclear Astrophysics: CEE, Michigan State University, East Lansing, Michigan 48824, USA

³Department of Physics and Astronomy, Michigan State University, East Lansing, Michigan 48824, USA

⁴Lawrence Berkeley National Laboratory, Berkeley, California 94720, USA

⁵Department of Physics and Astronomy, The University of North Carolina at Chapel Hill, Chapel Hill, North Carolina 27599, USA

⁶Department of Physics, Central Michigan University, Mt. Pleasant, Michigan 48859, USA

⁷Department of Physics and Astronomy, Swarthmore College, Swarthmore, Pennsylvania 19081, USA

⁸Department of Physics, Tohoku University, Sendai, Miyagi 980-8578, Japan



(Received 1 February 2019; revised manuscript received 1 April 2019; published 23 September 2019)

The Gamow-Teller strength distribution from ^{88}Sr was extracted from a $(t, ^3\text{He} + \gamma)$ experiment at 115 MeV/u to constrain estimates for the electron-capture rates on nuclei around $N = 50$, between and including ^{78}Ni and ^{88}Sr , which are important for the late evolution of core-collapse supernovae. The observed Gamow-Teller strength below an excitation energy of 8 MeV was consistent with zero and below 10 MeV amounted to 0.1 ± 0.05 . Except for a very-weak transition that could come from the 2.231-MeV 1^+ state, no γ lines that could be associated with the decay of known 1^+ states were identified. The derived electron-capture rate from the measured strength distribution is more than an order of magnitude smaller than rates based on the single-state approximation presently used in astrophysical simulations for most nuclei near $N = 50$. Rates based on shell-model and quasiparticle random-phase approximation calculations that account for Pauli-blocking and core-polarization effects provide better estimates than the single-state approximation, although a relatively strong transition to the first 1^+ state in ^{88}Rb is not observed in the data. Pauli-unblocking effects due to high stellar temperatures could partially counter the low electron-capture rates. The new data serve as a zero-temperature benchmark for constraining models used to estimate such effects.

DOI: [10.1103/PhysRevC.100.032801](https://doi.org/10.1103/PhysRevC.100.032801)

Introduction. Core-collapse supernovae (CCSNe) are among the most energetic explosions observed in the universe. They contribute to nucleosynthesis, stimulate galactic chemical evolution, and are birth places of neutron stars and black holes [1–4]. A very large fraction of the energy released in CCSNe is in the form of neutrinos, but the small fraction of energy released in the form of visible light is important for probing the mechanism of the explosion. In addition, CCSNe are predicted emission sites of gravitational waves [5,6]. Consequently, CCSNe are attractive sites for improving our understanding of the universe through multimessenger studies [7]. The accurate and detailed description of relevant nuclear physics processes is key to understanding the evolution of CCSNe and interpreting the multimessenger signals [8].

Nuclear-weak interaction processes, in particular, electron captures (ECs), are essential ingredients for simulating and understanding the dynamical evolution of the CCSNe [9–11]. EC reactions on nuclei in the upper pf and $pf\bar{g}/sd\bar{g}$ shells are particularly important during the collapse phase [12]. It was recently shown [13–16] that ECs on a group of about 75 nuclei around neutron number $N = 50$ between and including ^{78}Ni and ^{88}Sr (hereafter referred to as the high-sensitivity region) are responsible for about 50% of the uncertainties

in characteristic parameters, such as lepton fraction, entropy, mass enclosed at core bounce, and in-fall velocity [15]. Also, the EC rates on nuclei in this mass region could have a significant impact on the nucleosynthesis of trans-iron elements produced in thermonuclear supernovae [17].

EC rates are derived from Gamow-Teller (GT) transition-strength $[B(\text{GT})]$; here defined such that the strength associated with the decay of the free neutron has $B(\text{GT}) = 3]$ distributions in the β^+ direction. The EC rates presently used for the nuclei in the high-sensitivity region rely on an approximation that uses a single GT transition for which the strength and excitation energy were fitted to best reproduce EC rates for nuclei in the pf shell near stability [18,19]. This single-state approximation, which assumes a single strong Gamow-Teller transition with $B(\text{GT}) = 4.6$ and effective excitation energy adjusted based on the neutron and proton numbers of the parent nucleus [19], does not account for strong Pauli-blocking effects for heavier nuclei near $N = 50$, even for nuclei that are close to stability. The Pauli-blocking effects are caused by neutrons that occupy the nuclear orbits that would otherwise be available for proton-hole neutron-particle GT transitions in the β^+ direction. These effects could strongly reduce the EC rates for neutron-rich nuclei in the

high-sensitive region [13,15]. It is important to verify such effects experimentally and provide data to benchmark and guide theoretical calculations that are used to estimate the EC rates for the astrophysical simulations. At high stellar temperatures, Pauli unblockings are expected to become significant [18,20,21], and it is important that models used to estimate such effects are first validated at $T = 0$. However, for nuclei in which the Gamow-Teller transitions are not completely Pauli blocked, such as for ^{88}Sr , such temperature-dependent rate effects are expected to be relatively small [20].

The only way to experimentally extract GT strength distributions in the β^+ direction for neutron-rich nuclei is through the use of (n, p) -type charge-exchange (CE) reactions as the β^+ /EC-decay Q values for these nuclei are negative. CE experiments at intermediate beam energies ($\gtrsim 100$ MeV/u) provide an indirect method to extract the $B(\text{GT})$ distributions without Q -value constraints, based on a well-established proportionality between the CE cross section at zero momentum transfer and the $B(\text{GT})$ [22–24]. In this Rapid Communication, we present results of a $(t, {}^3\text{He} + \gamma)$ experiment aimed at extracting the GT transition strength [$B(\text{GT})$; associated with the transfer of $\Delta S = 1$ (spin), $\Delta T = 1$ (isospin), and $\Delta L = 0$ (angular momentum)] from the $N = 50, Z = 38$ nucleus ^{88}Sr , which is among the most proton-rich nuclei in the high-sensitive region. By combining the $(t, {}^3\text{He})$ CE reaction with high-resolution γ -ray detection, even weakly excited low-lying GT transitions that may relatively strongly impact the EC rates can be identified, achieving a sensitivity to states with GT strengths of as small as 0.01 [25,26]. This level of experimental sensitivity more or less coincides with the limit on the applicability of the use of charge-exchange reactions for reliably extracting GT strengths. This is due to interference effects between the central $\sigma\tau$ and the tensor- τ components of the nucleon-nucleon force. For strengths below 0.01, such effects complicate the clean identification of GT transitions from other transitions and introduce sizable (30% for GT strengths of about 0.01) systematic uncertainties [25,27,28].

The results discussed here are part of a broader effort to improve the electron-capture rates on nuclei in the high-sensitivity region. These efforts include additional experiments on other nuclei in this region, the incorporation of theoretical nuclear structure models aimed at improving the GT strength distributions used for electron-capture rate calculations, and astrophysical simulations similar to those in Refs. [13,15].

Experiment. A secondary triton beam was produced by fragmentation of a 150-MeV/u ^{16}O primary beam from the National Superconducting Cyclotron Laboratory (NSCL) Coupled Cyclotron Facility (CCF) on a 3525-mg/cm²-thick ^9Be production target placed at the entrance on the A1900 fragment separator [29]. A 99%-pure 115-MeV/u triton beam of 4×10^6 pps was generated with a momentum width of 0.5% (FWHM) by using a 195-mg/cm²-thick Al degrader in the A1900 intermediate image [30]. The beam was transported in a dispersion-matched mode [31,32] to an isotopically enriched ^{88}Sr (99.9%-pure) foil with a thickness of 19.6 mg/cm² placed at the S800 Spectrograph [33] pivot point. Due to the high reactivity of strontium, a special transfer system was used to insert the target without coming

into contact with air. ${}^3\text{He}$ ejectiles produced in the reaction were momentum analyzed and identified on the S800 focal plane [34]. The particle identification was performed on an event-by-event basis using the energy loss measured in a 5-mm-thick focal-plane scintillator and the time of flight relative to the CCF radio-frequency signal. Scattering angles and momenta of the ejectiles at the target location were reconstructed by ray tracing the angles and positions measured in two cathode-readout drift chambers on the S800 focal plane by using a fifth-order ion-optical inverse matrix calculated in COSY INFINITY [35]. Subsequently, the excitation energy (E_x) of the ^{88}Rb particles was determined in a missing-mass calculation up to 25 MeV with a resolution of 500 keV (FWHM), which is due to the intrinsic resolution that can be achieved in $(t, {}^3\text{He})$ experiments with a secondary triton beam and the difference in energy loss between tritons and ${}^3\text{He}$ particles in the ^{88}Sr target.

Scattering angles in the center-of-mass (c.m.) frame were measured in the range of $0^\circ < \theta_{\text{c.m.}} < 5.5^\circ$. A total luminosity of 2×10^{32} cm⁻² was achieved over 5 days. Data were acquired for the $^{12}\text{C}(t, {}^3\text{He})^{12}\text{B}(1^+, \text{g.s.})$ reaction by using a 2.6-mg/cm²-thick polystyrene $(\text{C}_8\text{H}_8)_n$ target. Its well-known cross section [24] was used to calibrate a nonintercepting primary-beam current probe that served as an absolute measure for the triton beam intensity during the ^{88}Sr runs.

The Gamma-Ray Energy Tracking In-beam Nuclear Array (GRETINA) [36,37], consisting of thirty-two 36-fold segmented high-purity Ge detectors mounted on a hemisphere and providing about 1π solid-angle coverage, was positioned around the ^{88}Sr target. The use of GRETINA allowed for the precise determination of γ -ray energies with a high photopeak-detection efficiency ($\sim 4\%$ at 2 MeV).

Experimental results. Double-differential cross sections for the $^{88}\text{Sr}(t, {}^3\text{He})$ reaction were generated in 0.5-MeV-wide bins in E_x . The average statistical error for each bin was 5%. The systematic error was $\sim 7\%$, dominated by the uncertainty in the triton beam intensity. Examples for $E_x = 2.25$ and 20.25 MeV are shown in Figs. 1(a) and 1(b), respectively. To extract the monopole contribution from the cross sections, a MDA [38,39] was performed for each bin in E_x by fitting the differential cross section with a linear combination of distorted-wave Born approximation (DWBA) angular distributions for angular momentum transfers of $\Delta L = 0$ –3. The DWBA calculations were performed using the code FOLD/DWHI [40]. The optical model potential (OMP) parameters were taken from Ref. [41]. Following Ref. [42], the depths of the OMP for the triton in the incoming channel were scaled by a factor of 0.85 from those for ${}^3\text{He}$ in the outgoing channel. The effective nucleon-nucleon interaction of Franey and Love [43] was double-folded over the transition densities of t - ${}^3\text{He}$ and ^{88}Sr - ^{88}Rb systems. The transition densities for t and ${}^3\text{He}$ were taken from variational Monte Carlo calculations [44]. For the ^{88}Sr - ^{88}Rb system, one-body transition densities were generated by using the shell-model (SM) code described below. Examples of MDA are shown in Figs. 1(a) and 1(b). The MDA results for $\theta_{\text{c.m.}} = 0.67^\circ$ and 1.56° as a function of E_x are shown in Figs. 1(c) and 1(d), respectively. For $E_x < 8$ MeV, the $\Delta L = 0$ contribution of the cross section is consistent with zero within the error bars (0.07 ± 0.1 mb/sr).

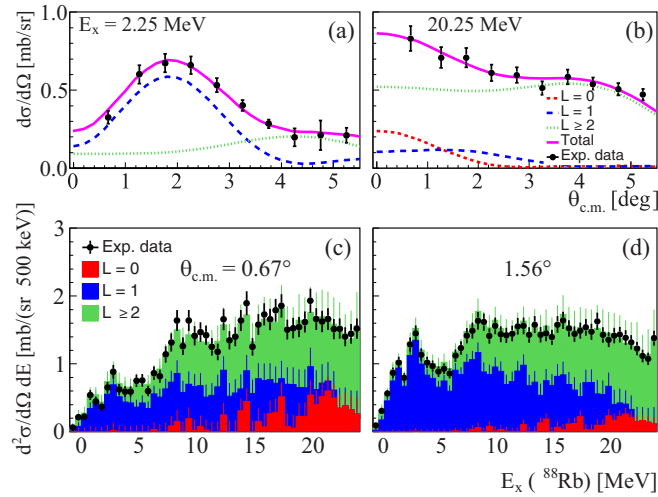


FIG. 1. Angular distributions for the $^{88}\text{Sr}(t, ^3\text{He})$ reaction at (a) $E_x = 2.25$ MeV and (b) 20.25 MeV, fitted in the multipole decomposition analysis (MDA). (Bottom) Double-differential cross sections for scattering angles in the ranges of (c) $0^\circ < \theta_{\text{c.m.}} < 1^\circ$ and (d) $1^\circ < \theta_{\text{c.m.}} < 2^\circ$. The colors represent contributions from excitations with different angular momentum transfers.

For $E_x > 10$ MeV, $\Delta L = 0$ contributions are observed, but the isovector spin-monopole resonance is expected to start contributing significantly in this region space [45].

The $B(\text{GT})$ strength was extracted from the $\Delta L = 0$ cross section at $\theta_{\text{c.m.}} = 0^\circ$ by using the proportionality relation: $\sigma_{L=0}(0^\circ) = \hat{\sigma}_{\text{GT}} F(q, \omega) B(\text{GT})$ [22–24]. $\hat{\sigma}_{\text{GT}}$ is the GT unit cross section, which was calculated (5.94 mb/sr) by the mass-dependent empirical relationship of Ref. [23], which has an uncertainty of 10%. $F(q, \omega)$ is a kinematic correction factor that depends on the momentum (q) and energy (ω) transfers and is obtained from DWBA calculations [22]. Its value was 1.2 (2.1) at $E_x = 0(10)$ MeV. Figure 2 shows the extracted

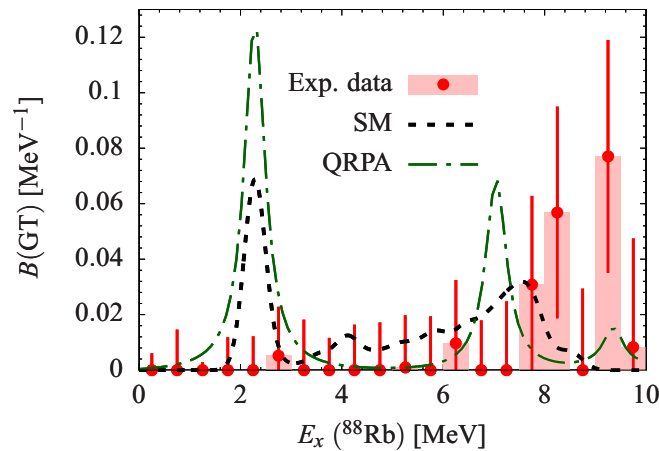


FIG. 2. $B(\text{GT})$ distribution extracted from MDA for $E_x < 10$ MeV. The error bars denote only the statistical uncertainties. The dashed lines correspond to theoretical calculations, SM, and quasiparticle random-phase approximation (QRPA) as described in the text.

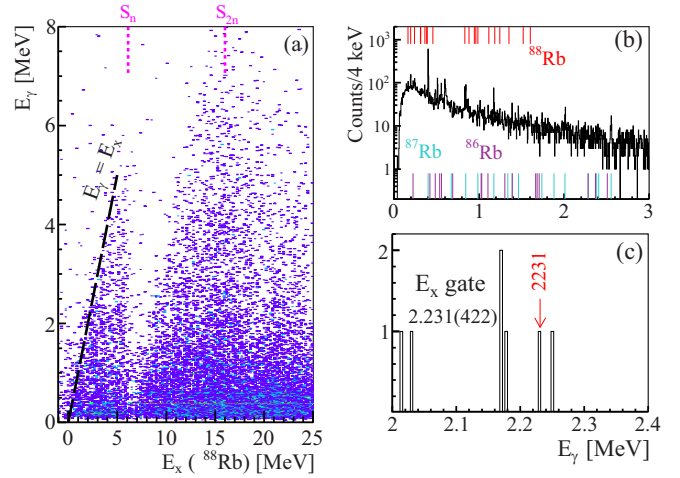


FIG. 3. (a) Two-dimensional histogram of γ -ray energy (E_γ) versus excitation energy (E_x) of ^{88}Rb . One- and two-neutron separation energies are indicated on top. (b) Projection onto the E_γ axis for $E_x \leq 25$ MeV. The color lines on the top and bottom of the figure represent the position of the observed transitions for $^{88,87,86}\text{Rb}$. (c) E_γ spectrum gated at $E_x = 2.231 \pm 0.422$ MeV.

$B(\text{GT})$ distribution in the energy range from 0 to 10 MeV. Above 10 MeV, the excitation of the isovector spin giant monopole resonance contributes to the monopole contributions in the excitation-energy spectra and becomes stronger than the excitation of Gamow-Teller strengths with increasing excitation energy [45–47]. Hence, only excitation energies below 10 MeV are considered here for the extraction of the GT strengths.

The only energy bin below the neutron separation energy with nonzero $B(\text{GT})$ is located at 2.75 ± 0.25 MeV, which correlates with the locations of several 1^+ states in ^{88}Rb known from the β^- decay of ^{88}Kr [48,49]. The summed $B(\text{GT})$ below $E_x = 10$ MeV is 0.10 ± 0.05 , although most of that comes from the region above 8 MeV. This result is significantly lower than the summed $B(\text{GT})$ of $0.7 \pm 0.1(\text{stat.}) \pm 0.1(\text{sys.})$ measured for the $^{90}\text{Zr}(n, p)$ reaction up to $E_x = 10$ MeV [50]. Based on transfer reaction experiments [51], the proton $0g_{9/2}$ occupation is 0.7 (1.0) for ^{88}Sr (^{90}Zr). Therefore, the decrease in GT strength observed below $E_x = 10$ MeV for ^{88}Sr as compared to ^{90}Zr is stronger than expected based on the proton $0g_{9/2}$ occupation number only.

Additional constraints on the GT strength can be obtained from the $(t, ^3\text{He} + \gamma)$ coincident data. Figure 3(a) shows the two-dimensional histogram that correlates the energy of γ rays (E_γ) with $E_x(^{88}\text{Rb})$. Due to the wide E_x range covered, γ -ray transitions from states in ^{88}Rb , ^{87}Rb , and ^{86}Rb were observed as shown in Fig. 3(b). The nonobservation of γ rays from ^{87}Kr indicates that the probability of decay by proton emission from ^{88}Rb was very small.

By setting narrow gates on E_x determined from the $(t, ^3\text{He})$ reaction, the γ spectrum for low $E_x(^{88}\text{Rb})$ was investigated for evidence for the decay from known 1^+ states or for unknown γ lines that could stem from previously unknown 1^+ states. No significant signals were found with the exception of the observation of a single event that could be due to decay from

the known 2.231-MeV 1^+ state as shown in Fig. 3(c). This spectrum was obtained by setting a gate on $E_x = 2.231 \pm 0.422$ MeV in the $^{88}\text{Sr}(t, ^3\text{He})$ spectrum where the width of the gate corresponds to 2σ of the energy resolution. By using a Bayesian analysis [52], it was determined that, with an 86% probability, the credible interval for $B(\text{GT})$ for the 2.231-MeV state ranges from 0 to 0.022, which includes the possibility that the observed count is not due to the decay from this state. The extracted Gamow-Teller strength from the MDA analysis in the relevant excitation energy bin for this transition is $0.006^{+0.02}_{-0.006}$.

Comparison with theory. The experimental results were compared to SM and QRPA calculations. The shell-model calculations, performed with the code NUSHELLX [53], assumed a ^{78}Ni core and a valence space of $(0f_{5/2}, 1p_{3/2}, 1p_{1/2}, 0g_{9/2})$ for protons and $(0g_{7/2}, 1d_{5/2}, 1d_{3/2}, 2s_{1/2}, 0h_{11/2})$ for neutrons. The proton-proton and proton-neutron two-body matrix elements were obtained from the J44PNA effective interaction [54] and a renormalized G matrix using the charge-dependent-Bonn nucleon-nucleon interaction [55], respectively.

The single-particle energies were determined from the observed single-particle states in ^{89}Sr . To account for the model-space truncation, the result of the calculation was scaled by a factor $\frac{1}{h}$, where h is a hindrance factor that is a product of two factors: h_{high} and $h_{\text{c.p.}}$ [56]. h_{high} is associated with the admixtures of two-particle two-hole states with unperturbed energies of $2\hbar\omega$ and higher in the oscillator basis. This factor accounts for the well-known quenching of the GT transition strength [57,58]. The empirical value for the pf model space $h_{\text{high}} = 1.81$ [59] was used. $h_{\text{c.p.}}$ is due to the core polarization for the $0g$ orbital. It accounts for the mixing between $0g_{9/2}$ and $0g_{7/2}$ spin-orbit partners and depends on the proton occupation number in the $0g_{9/2}$ orbital. $h_{\text{c.p.}}$ is largest when the number of $0g_{9/2}$ protons is small [56]. An occupation number of 0.58 was calculated for the $\pi 0g_{9/2}$ shell in ^{88}Sr by using the Ji/Wildenthal effective interaction [60], which is close to the experimental value of 0.7 [51]. The hindrance due to the core polarization was taken from the results of Towner in Ref. [56] (Table 5). The value for two protons in $0g_{9/2}$ of $h_{\text{c.p.}} = 5.0$, obtained from the $\pi + \rho$ interaction (the range for the three interactions given is 3.5–5.9), was used in our calculation. The $0g_{9/2}$ proton number dependence of the hindrance factor $h_{\text{high}} \times h_{\text{c.p.}}$ leads to a Z -dependent hindrance factor that is consistent with that deduced from the β^+ decay of nuclei with $N = 50$ ranging from ^{94}Ru up to ^{100}Sn [61–64].

The QRPA calculation was performed by using the axially deformed Skyrme finite amplitude method [65,66]. This method has recently been extended to odd- A nuclei in the equal-filling approximation [67] and is, therefore, a candidate for calculating GT strengths and EC rates for a large number of nuclei and replacing the EC rates based on the single-state approximation discussed above. The Skyrme functional and single-particle space model are the same as those used in the global calculation of Ref. [68], which fixed a single set of parameters including an effective axial-vector coupling constant g_A of 1.0.

The theoretical calculations shown in Fig. 2 have been folded with the experimental resolution and the excitation energy of the first 1^+ state was matched to that of the first

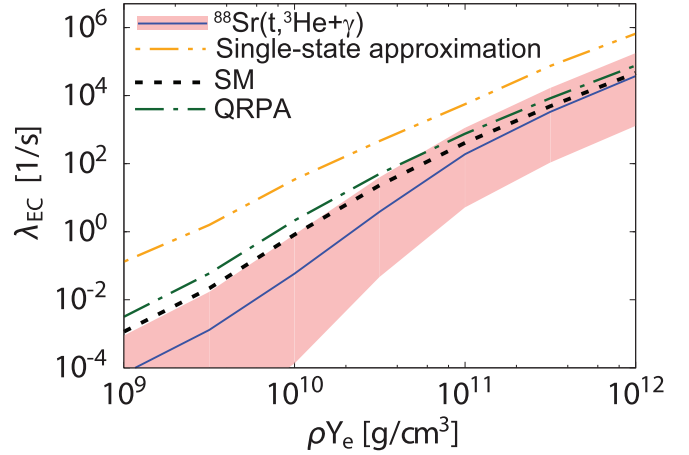


FIG. 4. EC rates on ^{88}Sr as a function of stellar density at a temperature of 10^{10} K. The shaded band with the solid central curve represents the result based on the $^{88}\text{Sr}(t, ^3\text{He} + \gamma)$ data. The dashed and dot-dashed curves are based on the SM and QRPA calculations, respectively. The dot-dot-dashed line represents the approximate method for estimating the EC rate.

known 1^+ state in ^{88}Rb (at $E_x = 2.231$ MeV). The SM and QRPA calculations both predict a strong transition to the first 1^+ state that is not observed experimentally. The summed strength up to $E_x = 10$ MeV is 0.12 (0.14) for the SM (QRPA) calculations. These summed values are consistent with the present data of 0.1 ± 0.05 . The results indicate that Pauli blocking and structural effects (core polarization) play an important role in the reduction of the Gamow-Teller strength at low excitation energies. In addition, the results suggest that the Gamow-Teller strength is distributed over a wider excitation-energy range than the calculations predict.

Electron-capture rates. Stellar EC rates (λ_{EC}) were calculated based on the formalism in Refs. [69–72] in a code previously used in Refs. [25,26,73,74]. Only transitions from the ground state of ^{88}Sr were considered here. Figure 4 shows the calculated EC rates (based on the experimental and theoretical GT strength distributions) during the late stages of CCSN, just prior to the bounce, during which the stellar density ranges from 10^9 to 10^{12} g/cm 3 and the temperature is $\sim 10^{10}$ K. Because no known 1^+ state exists below $E_x = 2$ MeV and the MDA analysis found no indication for any GT strength up to that energy, the first transition assumed to contribute to the EC rate based on the data was the 2.231-MeV state, with an upper limit to the strength based on the γ -decay analysis [$B_{\text{up}}(\text{GT}) = 0.022$]. The Q value for EC on ^{88}Sr is -4.8 MeV, which means that, only at a density of 10^{11} g/cm 3 , the Fermi energy of ~ 15 MeV is sufficiently high to cover the strength distribution up to $E_x = 10$ MeV and that details of the GT strength distribution below that E_x matter up to that density. The higher the density, the less sensitive the EC rate to details of the strength distribution as most of the strength distribution is below the Fermi energy.

Due to the presence of the relatively strong transition to the first 1^+ state, the EC rates based on the QRPA and SM calculations are higher than the upper limit set by the data for stellar densities below $\sim 10^{10}$ g/cm 3 . At higher densities,

the rates based on the SM and QRPA calculations are within the upper limit set by the data since the summed strengths up to 10 MeV are, within error bars, consistent. The EC rates calculated based on the single-state approximation are more than an order of magnitude too high.

Considering that ^{88}Sr is the among the most proton-rich $N = 50$ nucleus in the high-sensitivity region (with the least Pauli blocking), it is very likely that the rates based on the approximation will be also be much too high for the other nuclei in the high-sensitivity region. This has a strong impact on the dynamical evolution during the collapse phase [13,15]. The drop in lepton fraction during the collapse reduces by 10%, and the enclosed mass at core bounce increases by 10% when the EC rates for nuclei in the high-sensitivity region are reduced by a factor of 10. Finally, we note that, although Pauli-unblocking effects due to the high temperature in the collapsing star should increase the EC rates compared to rates shown in Fig. 4, the structural (core-polarization) effects are equally important, especially for nuclei, such as ^{88}Sr in which Pauli blocking is not complete at $T = 0$, and must be considered in theoretical models used for estimating EC rates at elevated temperatures. Therefore, the present data provide an important zero-temperature benchmark for such theoretical estimates.

Clearly, the further development of theoretical models is important. For the shell-model calculations, it will be key to increase the model space to include, at least, the $g_{9/2}$ and $g_{7/2}$ orbits for both protons and neutrons and to have an appropriately renormalized Hamiltonian for that model space. Calculations in this larger model space have recently been performed for Zr isotopes [75] and in the future can be tested for calculations of GT transitions. To improve on the QRPA calculations, one has to include multiquasiparticle excitations, for example, following the techniques described in recent works from Refs. [76–78]. For applications in astrophysical

modeling, it is important that such calculations can be performed for a wide variety of nuclei.

Summary. The GT transition strength in ^{88}Sr was measured in a high-resolution ($t, {}^3\text{He} + \gamma$) experiment to gain insight in EC rates of nuclei near $N = 50$ above ^{78}Ni that are most important during the collapse phase of massive stars prior to the supernova explosion. The extracted $B(\text{GT})$ is consistent with zero in the energy range from 0 up to $E_x = 8$ MeV and sums to 0.1 ± 0.05 up to $E_x = 10$ MeV. SM and QRPA calculations are consistent with this summed strength but predict a relatively strong transition to a low-lying state not observed in the experiment. As the most proton-rich $N = 50$ nucleus in the high-sensitivity region, these results indicate that the EC rates based on a single-state approximation that is used in astrophysical simulations are too high. Although Pauli-unblocking effects due to the high stellar temperatures during the collapse phase partially counter the lowering of the EC rates, the results show that structural effects must be carefully considered as they significantly lower the GT strengths and EC rates. Hence, the present data also serve as a zero-temperature benchmark for theoretical models that can be used to estimate temperature-dependent Pauli-unblocking effects.

Acknowledgments. We thank the NSCL staff for their support. This work was supported by the U.S. National Science Foundation (NSF) under Cooperative Agreements No. PHY-156554 (NSCL), No. PHY-1430152 (JINA Center for the Evolution of the Elements), and No. PHY-1811855. GREINA was funded by the U.S. Department of Energy in the Office of Nuclear Physics of the Office of Science. Operation of the array at NSCL was supported by the DOE under Grants No. DE-SC0014537 (NSCL) and No. DE-AC02-05CH11231 (LBNL). The ^{88}Sr foils used in this research were supplied by the U.S. Department of Energy Office of Science by the Isotope Program in the Office of Nuclear Physics.

-
- [1] C. L. Fryer, *Astrophys. J.* **522**, 413 (1999).
 - [2] A. Heger, C. L. Fryer, S. E. Woosley, N. Langer, and D. H. Hartmann, *Astrophys. J.* **591**, 288 (2003).
 - [3] H.-T. Janka, K. Langanke, A. Marek, G. Martínez-Pinedo, and B. Müller, *Phys. Rep.* **442**, 38 (2007).
 - [4] A. Burrows, *Rev. Mod. Phys.* **85**, 245 (2013).
 - [5] H. Dimmelmeier, C. D. Ott, A. Marek, and H.-T. Janka, *Phys. Rev. D* **78**, 064056 (2008).
 - [6] S. Richers, C. D. Ott, E. Abdikamalov, E. O'Connor, and C. Sullivan, *Phys. Rev. D* **95**, 063019 (2017).
 - [7] K. Nakamura *et al.*, *Mon. Not. R. Astron. Soc.* **461**, 3296 (2016).
 - [8] A. Arcones *et al.*, *Prog. Part. Nucl. Phys.* **94**, 1 (2017).
 - [9] H. Bethe, G. Brown, J. Applegate, and J. Lattimer, *Nucl. Phys.* **A324**, 487 (1979).
 - [10] K. Langanke and G. Martínez-Pinedo, *Rev. Mod. Phys.* **75**, 819 (2003).
 - [11] T. Suzuki, H. Toki, and K. Nomoto, *Astrophys. J.* **817**, 163 (2016).
 - [12] W. R. Hix, O. E. B. Messer, A. Mezzacappa, M. Liebendörfer, J. Sampaio, K. Langanke, D. J. Dean, and G. Martínez-Pinedo, *Phys. Rev. Lett.* **91**, 201102 (2003).
 - [13] C. Sullivan, E. O'Connor, R. G. T. Zegers, T. Grubb, and S. M. Austin, *Astrophys. J.* **816**, 44 (2016).
 - [14] S. Furusawa, H. Nagakura, K. Sumiyoshi, C. Kato, and S. Yamada, *Phys. Rev. C* **95**, 025809 (2017).
 - [15] R. Titus, C. Sullivan, R. G. T. Zegers, B. A. Brown, and B. Gao, *J. Phys. G: Nucl. Part. Phys.* **45**, 014004 (2018).
 - [16] A. Pascal *et al.*, Impact of electron capture rates on nuclei far from stability on core-collapse supernovae, [arXiv:1906.05114](https://arxiv.org/abs/1906.05114).
 - [17] S. Jones *et al.*, *Astron. Astrophys.* **622**, A74 (2019).
 - [18] K. Langanke *et al.*, *Phys. Rev. Lett.* **90**, 241102 (2003).
 - [19] A. R. Raduta, F. Gulminelli, and M. Oertel, *Phys. Rev. C* **95**, 025805 (2017).
 - [20] K. Langanke, E. Kolbe, and D. J. Dean, *Phys. Rev. C* **63**, 032801(R) (2001).
 - [21] A. A. Dzhiyev, A. I. Vdovin, V. Y. Ponomarev, J. Wambach, K. Langanke, and G. Martínez-Pinedo, *Phys. Rev. C* **81**, 015804 (2010).
 - [22] T. Taddeucci *et al.*, *Nucl. Phys.* **A469**, 125 (1987).
 - [23] R. G. T. Zegers *et al.*, *Phys. Rev. Lett.* **99**, 202501 (2007).
 - [24] G. Perdikakis *et al.*, *Phys. Rev. C* **83**, 054614 (2011).
 - [25] S. Noji *et al.*, *Phys. Rev. Lett.* **112**, 252501 (2014).
 - [26] S. Noji *et al.*, *Phys. Rev. C* **92**, 024312 (2015).

- [27] R. G. T. Zegers *et al.*, *Phys. Rev. C* **74**, 024309 (2006).
- [28] G. W. Hitt *et al.*, *Phys. Rev. C* **80**, 014313 (2009).
- [29] D. Morrissey, B. Sherrill, M. Steiner, A. Stolz, and I. Wiedenhoever, *Nucl. Instrum. Methods Phys. Res., Sect. B* **204**, 90 (2003).
- [30] G. Hitt *et al.*, *Nucl. Instrum. Methods Phys. Res., Sect. A* **566**, 264 (2006).
- [31] B. Sherrill *et al.*, *Nucl. Instrum. Methods Phys. Res., Sect. A* **432**, 299 (1999).
- [32] M. Berz, K. Joh, J. A. Nolen, B. M. Sherrill, and A. F. Zeller, *Phys. Rev. C* **47**, 537 (1993).
- [33] D. Bazin, J. Caggiano, B. Sherrill, J. Yurkon, and A. Zeller, *Nucl. Instrum. Methods Phys. Res., Sect. B* **204**, 629 (2003).
- [34] J. Yurkon *et al.*, *Nucl. Instrum. Methods Phys. Res., Sect. A* **422**, 291 (1999).
- [35] K. Makino and M. Berz, *Nucl. Instrum. Methods Phys. Res., Sect. A* **427**, 338 (1999).
- [36] S. Paschalis *et al.*, *Nucl. Instrum. Methods Phys. Res., Sect. A* **709**, 44 (2013).
- [37] D. Weisshaar *et al.*, *Nucl. Instrum. Methods Phys. Res., Sect. A* **847**, 187 (2017).
- [38] B. Bonin *et al.*, *Nucl. Phys. A* **430**, 349 (1984).
- [39] M. Ichimura, H. Sakai, and T. Wakasa, *Prog. Part. Nucl. Phys.* **56**, 446 (2006).
- [40] J. Cook and J. A. Carr, computer program FOLD, Florida State University, (unpublished); based on F. Petrovich and D. Stanley, *Nucl. Phys. A* **275**, 487 (1977); modified as described in J. Cook, K. W. Kemper, P. V. Drumm, L. K. Fifield, M. A. C. Hotchkis, T. R. Ophel, and C. L. Woods, *Phys. Rev. C* **30**, 1538 (1984); R. G. T. Zegers, S. Fracasso, and G. Colò (unpublished).
- [41] J. Kamiya *et al.*, *Phys. Rev. C* **67**, 064612 (2003).
- [42] S. V. D. Werf *et al.*, *Nucl. Phys. A* **496**, 305 (1989).
- [43] M. A. Franey and W. G. Love, *Phys. Rev. C* **31**, 488 (1985).
- [44] S. C. Pieper and R. B. Wiringa, *Annu. Rev. Nucl. Part. Sci.* **51**, 53 (2001).
- [45] K. Miki *et al.*, *Phys. Rev. Lett.* **108**, 262503 (2012).
- [46] I. Hamamoto and H. Sagawa, *Phys. Rev. C* **62**, 024319 (2000).
- [47] C. J. Guess *et al.*, *Phys. Rev. C* **83**, 064318 (2011).
- [48] R. L. Bunting, W. L. Talbert, J. R. McConnell, and R. A. Meyer, *Phys. Rev. C* **13**, 1577 (1976).
- [49] E. McCutchan and A. Sonzogni, *Nucl. Data Sheets* **115**, 135 (2014).
- [50] K. Yako *et al.*, *Phys. Lett. B* **615**, 193 (2005).
- [51] A. Pfeiffer *et al.*, *Nucl. Phys. A* **455**, 381 (1986).
- [52] J. Ghosh, M. Delampady, and T. Samanta, *An Introduction to Bayesian Analysis: Theory and Methods*, Springer Texts in Statistics (Springer, New York, 2007).
- [53] B. A. Brown and W. D. M. Rae, *Nuclear Data Sheets* **120**, 115 (2014).
- [54] A. F. Lisetskiy, B. A. Brown, M. Horoi, and H. Grawe, *Phys. Rev. C* **70**, 044314 (2004).
- [55] I. Dillmann *et al.*, *Phys. Rev. Lett.* **91**, 162503 (2003).
- [56] I. Towner, *Nucl. Phys. A* **444**, 402 (1985).
- [57] C. Gaarde *et al.*, *Nucl. Phys. A* **369**, 258 (1981).
- [58] C. Gaarde, in *Proceedings of the Niels Bohr Centennial Conference on Nuclear Structure, Copenhagen*, edited by R. A. Broglia, G. B. Hagemann, and B. Herskind (North-Holland, Amsterdam, 1985), p. 449c.
- [59] G. Martínez-Pinedo, A. Poves, E. Caurier, and A. P. Zuker, *Phys. Rev. C* **53**, R2602(R) (1996).
- [60] X. Ji and B. H. Wildenthal, *Phys. Rev. C* **37**, 1256 (1988).
- [61] P. Graf and H. Münzel, *Radiochim. Acta* **20**, 140 (1973).
- [62] K. Rykaczewski *et al.*, *Z. Phys. A: At. Nucl.* **322**, 263 (1985).
- [63] A. Płochocki *et al.*, *Z. Phys. A: Hadrons Nucl.* **342**, 43 (1992).
- [64] C. B. Hinke *et al.*, *Nature (London)* **486**, 341 (2012).
- [65] P. Avogadro and T. Nakatsukasa, *Phys. Rev. C* **84**, 014314 (2011).
- [66] M. T. Mustonen, T. Shafer, Z. Zenginerler, and J. Engel, *Phys. Rev. C* **90**, 024308 (2014).
- [67] T. Shafer, J. Engel, C. Fröhlich, G. C. McLaughlin, M. Mumpower, and R. Surman, *Phys. Rev. C* **94**, 055802 (2016).
- [68] M. T. Mustonen and J. Engel, *Phys. Rev. C* **93**, 014304 (2016).
- [69] G. M. Fuller, W. A. Fowler, and M. J. Newman, *Astrophys. J., Suppl. Ser.* **42**, 447 (1980).
- [70] G. M. Fuller, W. A. Fowler, and M. J. Newman, *Astrophys. J.* **252**, 715 (1982).
- [71] G. M. Fuller, W. A. Fowler, and M. J. Newman, *Astrophys. J., Suppl. Ser.* **48**, 279 (1982).
- [72] G. M. Fuller, W. A. Fowler, and M. J. Newman, *Astrophys. J.* **293**, 1 (1985).
- [73] S. Gupta, E. F. Brown, H. Schatz, P. Möller, and K.-L. Kratz, *Astrophys. J.* **662**, 1188 (2007).
- [74] A. L. Cole, T. S. Anderson, R. G. T. Zegers, S. M. Austin, B. A. Brown, L. Valdez, S. Gupta, G. W. Hitt, and O. Fawwaz, *Phys. Rev. C* **86**, 015809 (2012).
- [75] T. Togashi, Y. Tsunoda, T. Otsuka, and N. Shimizu, *Phys. Rev. Lett.* **117**, 172502 (2016).
- [76] Y. F. Niu, G. Colò, E. Vigezzi, C. L. Bai, and H. Sagawa, *Phys. Rev. C* **94**, 064328 (2016).
- [77] Y. Niu, Z. Niu, G. Colò, and E. Vigezzi, *Phys. Lett. B* **780**, 325 (2018).
- [78] C. Robin and E. Litvinova, *Phys. Rev. C* **98**, 051301(R) (2018).